

Optimal Weights and Multiple Tracers in Large-Scale Structure

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NH, Uroš Seljak & Vincent Desjacques 2011; 1104.2321 (astro-ph.CO)

NH, Uroš Seljak, Vincent Desjacques, Robert E. Smith & Tobias Baldauf 2010; 1004.5377 (astro-ph.CO) + PRD

Uroš Seljak, NH & Vincent Desjacques 2009; 0904.2963 (astro-ph.CO) + PRL

Cosmological Non-Gaussianity: Observations Confront Theory Workshop
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Consider multiple different-mass tracers of the underlying dark matter density field:

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Each of them is sensitive to primordial non-Gaussianity through their linear bias:

$$b_i(k, f_{\text{NL}}) = b_i(f_{\text{NL}}=0) + f_{\text{NL}} [b_i(f_{\text{NL}}=0) - 1] u(\mathbf{k}, z)$$

with

$$u(\mathbf{k}, z) \equiv 3\delta_c \Omega_m H_0^2 / \mathbf{k}^2 T(\mathbf{k}) D(z) c^2$$

The 2-point clustering statistics is described by the halo covariance matrix in Fourier space:

$$C_{ij} \equiv \langle \delta_i \delta_j \rangle = b_i b_j P + \mathcal{E}_{ij}$$

with $P \equiv \langle \delta^2 \rangle$, $b_i \equiv \langle \delta_i \delta \rangle / \langle \delta^2 \rangle$, $\mathcal{E}_{ij} \equiv \langle \epsilon_i \epsilon_j \rangle$

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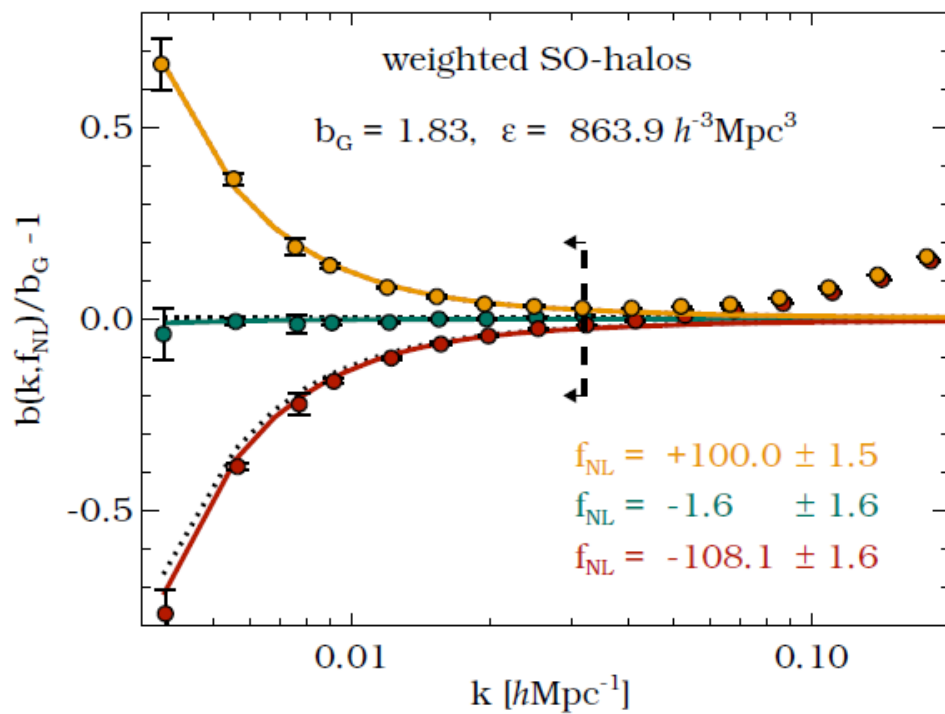
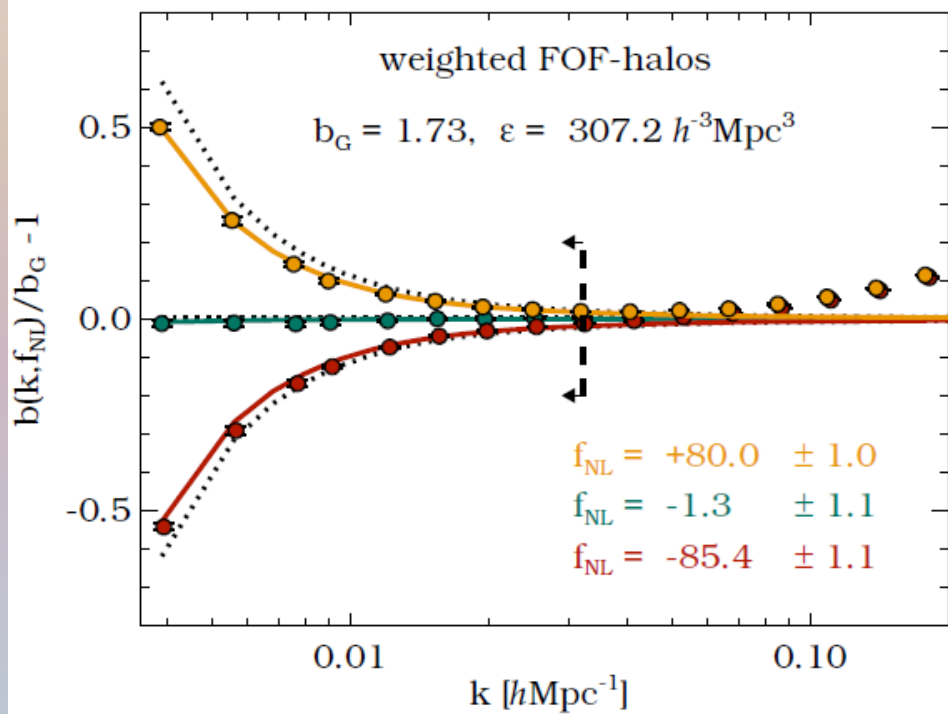
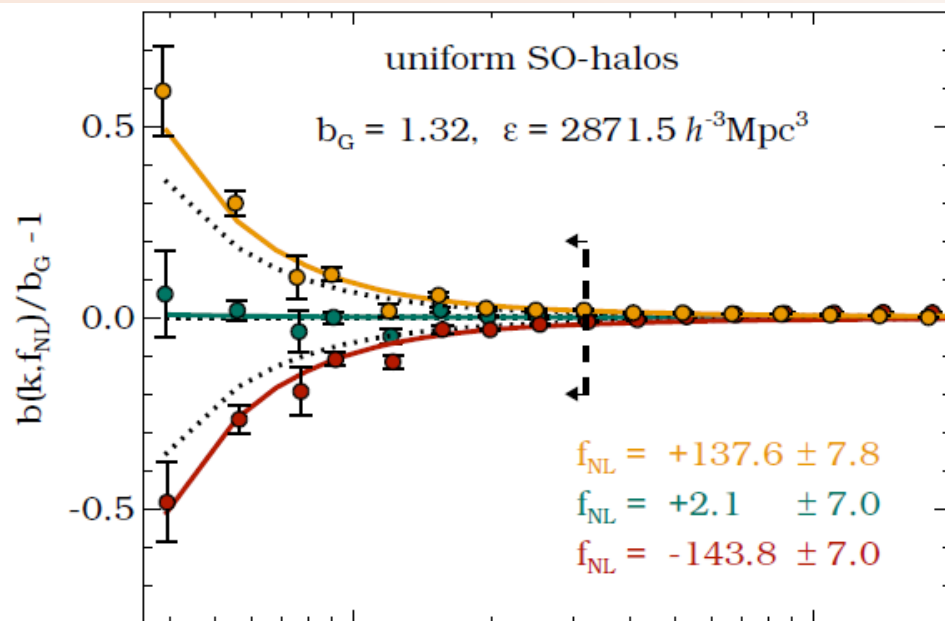
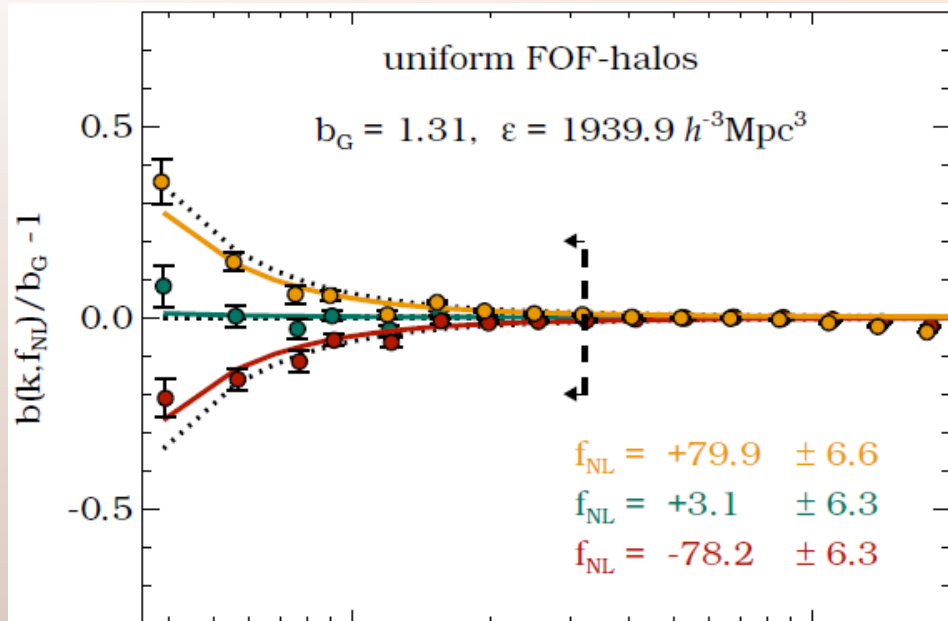
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With knowledge of δ (from simulations), we can

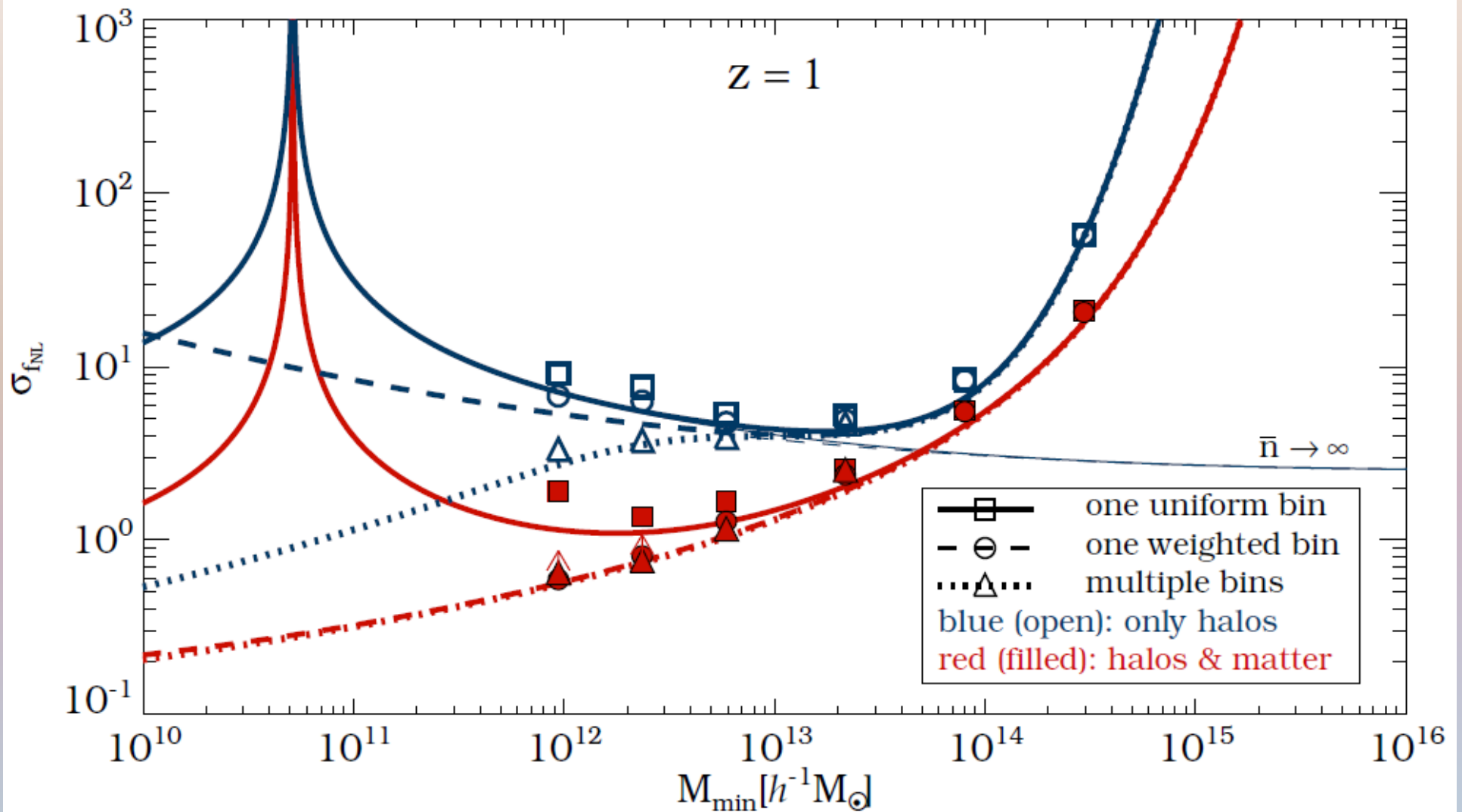
determine $\mathcal{E}_{ij} = \langle \delta_i \delta_j \rangle - \langle \delta_i \delta \rangle \langle \delta_j \delta \rangle / \langle \delta^2 \rangle$

- $\mathcal{E}_{ij} \neq \delta_{ij}^K / n_i$ (Poisson sampling)
- \mathcal{E}_{ij} exhibits one low eigenvalue $\lambda_- < 1/n_i$
- Eigenvector V_- provides optimal weighting

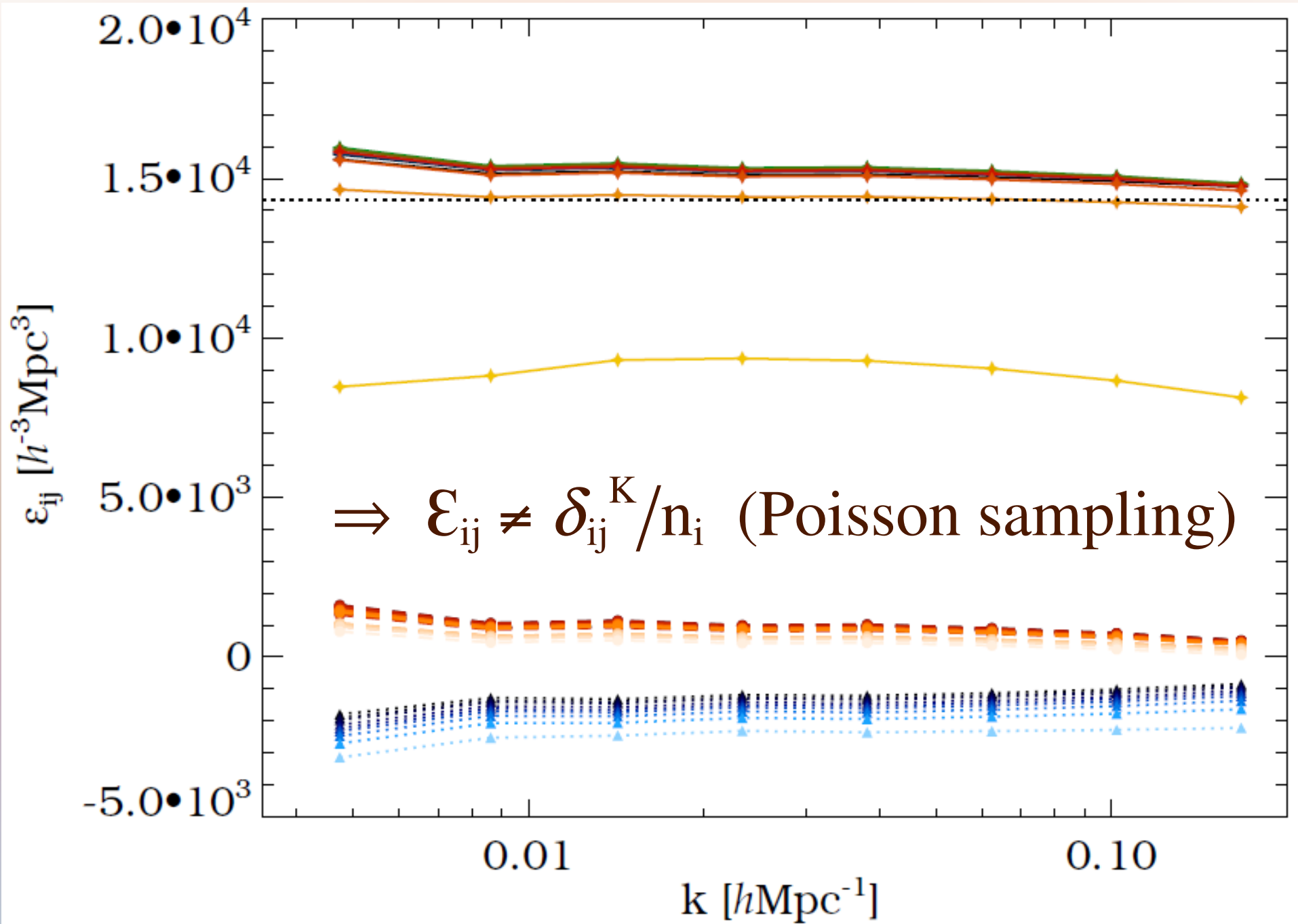


Fisher information on primordial non-Gaussianity:

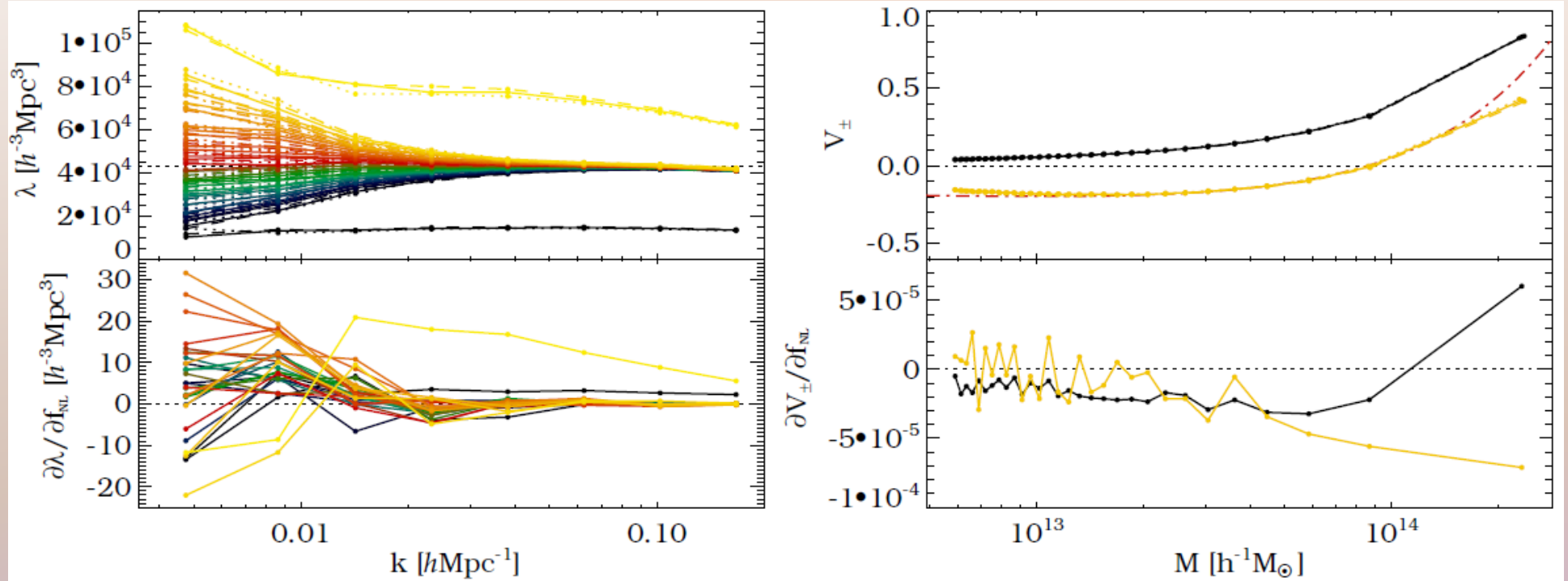
$$F_{f_{\text{NL}}f_{\text{NL}}} = \frac{1}{2} \text{Tr}[\mathbf{C}^{-1} \partial\mathbf{C}/\partial f_{\text{NL}} \mathbf{C}^{-1} \partial\mathbf{C}/\partial f_{\text{NL}}]$$



Thank you!



Eigendecomposition of Ξ_{ij} :



- Enhanced and suppressed eigenvalue w.r.t. $1/n_i$
- Eigenvectors provide weights for each mass bin

